

# LIGO: Laser Detection of Ripples in Space

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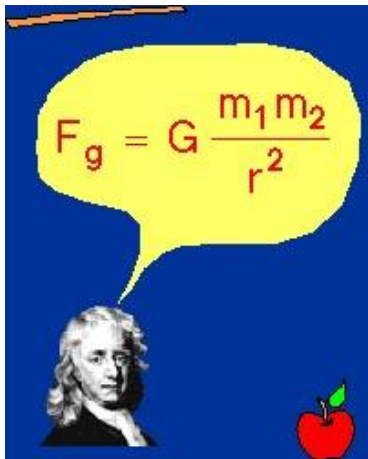
The College of William & Mary



May 25, 2013

- 1 History of gravity
  - Newton's laws
  - Einstein's laws
  - A bit of astrophysics
- 2 Detectors
  - Gravitational wave interferometer
- 3 Assorted LIGO pictures
- 4 other detectors
- 5 Extra information

# Newton's laws 1687

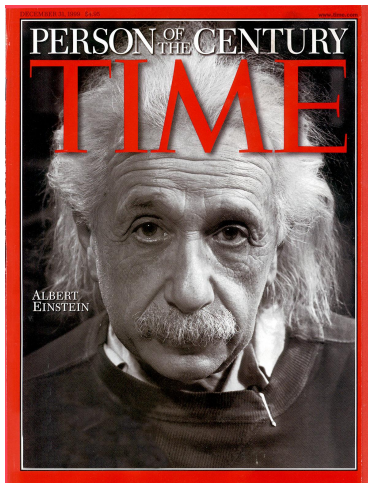


Laws of motion and law of gravitation solved problems of astronomy and terrestrial physics.

- eccentric orbits
- tides
- perturbation of moon orbit due to sun

Unified the work of Galileo, Copernicus and Kepler.

Did not explained precession of Mercury orbit



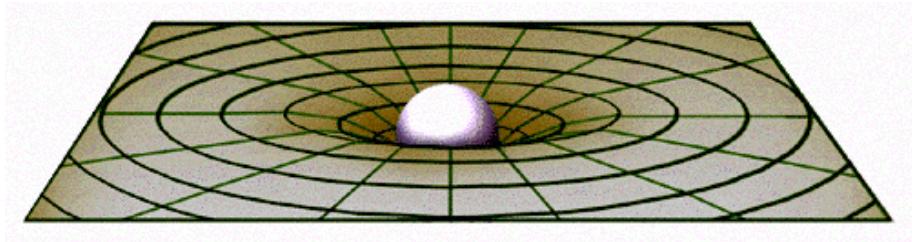
## The General Theory of Relativity and theory of Gravity (1916)

- No absolute motion  
thus only relative motion
- Space and time are not separate  
thus four dimensional space-time
- Gravity is not a force acting at a distance  
thus warpage of space-time



# General relativity

- A geometric theory connecting matter to spacetime
- Matter tells spacetime how to curve
- Spacetime tells matter how to move



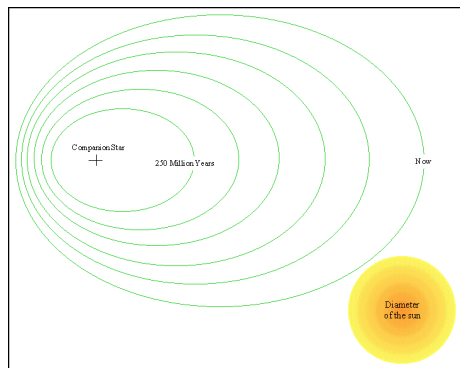
important predictions

- Light path bends in vicinity of massive object → confirmed in 1919
- Gravitational radiation (waves) → confirmed **indirectly** in 1974

# Indirect observation of gravitational wave

Emission of gravitational radiation from pulsar PSR1913+16 leads to loss of orbital energy

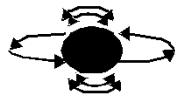
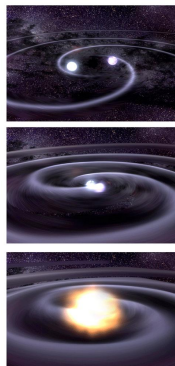
- orbital period decreased by 14 sec from 1975 to 1994
- measured to 50 msec accuracy
- deviation grows quadratically with time



Nobel prize in 1997 Taylor and Hulse

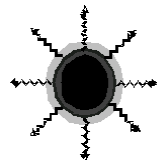
# Astrophysical sources of GW

- Coalescing compact binaries
  - objects: NS-NS, BH-NS, BH-BH
  - physics regimes: Inspiral, merger, ringdown
  
- Periodic sources
  - spinning neutron stars (pulsars)

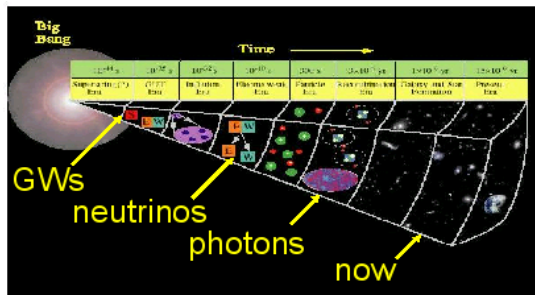


# Astrophysical sources of GW (cont)

- Burst events
  - Supernovae with asymmetric collapse



- Stochastic background
  - right after Big Bang ( $t = 10^{-43}$  sec)
  - continuum of sources



# Astrophysics with GWs vs. E&M

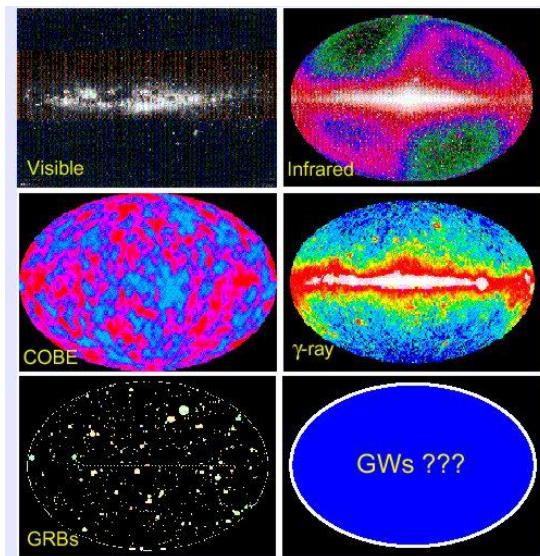
## E&M (photons)

- Space as medium for field
- Accelerating charge
- Absorbed, scattered, dispersed by matter
- 10 MHz and up
- Light = not dark (but >95% of Universe is dark)

## GW

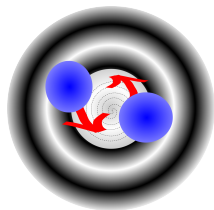
- Spacetime itself ripples
- Accelerating aspherical mass
- Very small interaction; matter is transparent
- 10 kHz and down
- Radiated by dark mass distributions

# New view to the universe

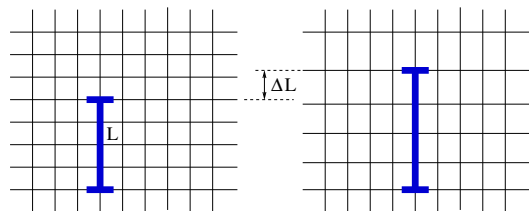


# Gravitational waves (GW)

- Predicted by the General Theory of Relativity
- Generated by aspherical mass distribution
- Induce space-time ripples which propagate with speed of light
- New tool for astrophysics



GW stretch and squeeze space-time thus move freely floating objects



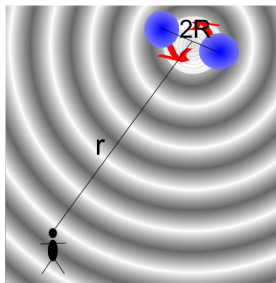
Strain - strength of GW

$$h = \frac{\Delta L}{L} \quad (1)$$

typical strain

$$h \sim 10^{-21} \quad (2)$$

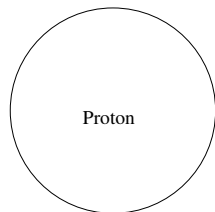
# Typical strain



## Two neutron star

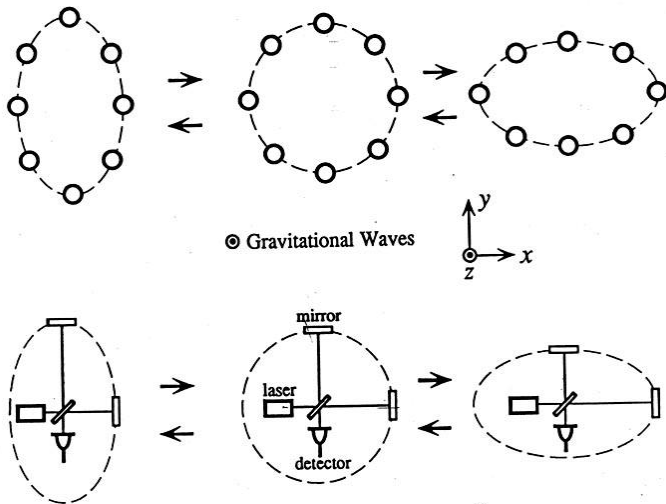
with a mass of 1.4 solar masses each  
orbiting each other with a frequency  $f = 400$  Hz  
at a distance  $2R = 20$  km  
would generate strain  $h \sim 10^{-21}$   
at distance equal to  $10^{23}$  m  
(distance to the Virgo cluster)  
For 4 km base line that would correspond to  
 $\Delta L$  thousand times smaller than size of proton.

Detection of GW is difficult problem

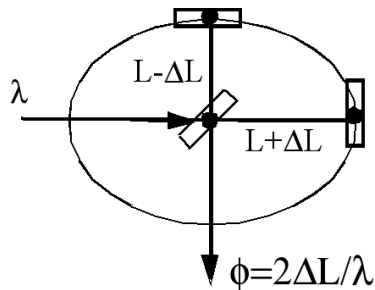
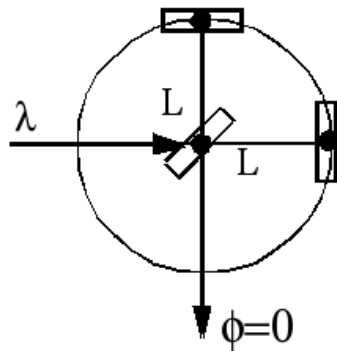




# GW acting on matter



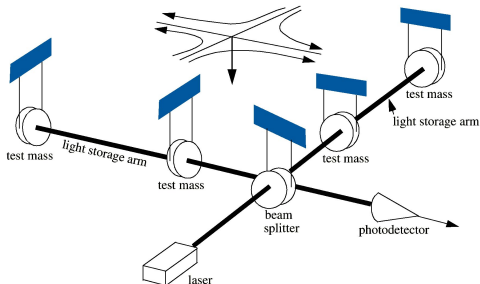
# Interferometric Measurement



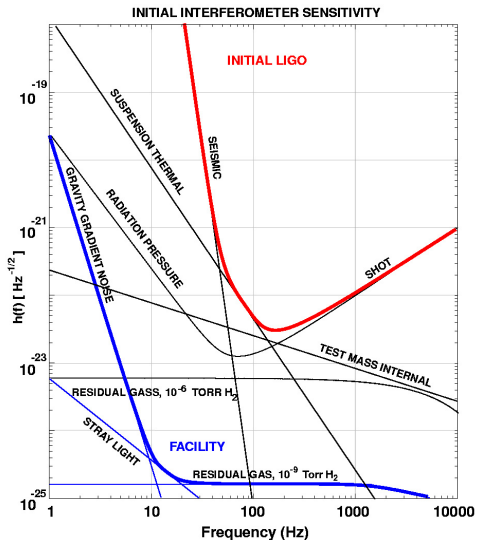
# Laser Interferometer Gravitational-wave Observatory



- $L = 4 \text{ km}$
- $h \sim 10^{-21}$
- $\Delta L \sim 10^{-18} \text{ m}$
- $\Delta \phi \sim 10^{-10} \text{ rad}$



# LIGO sensitivity goal and noise budget



## Displacement noise

- seismic
- thermal suspension
- thermal Brownian
- radiation pressure noise

## Detection noise

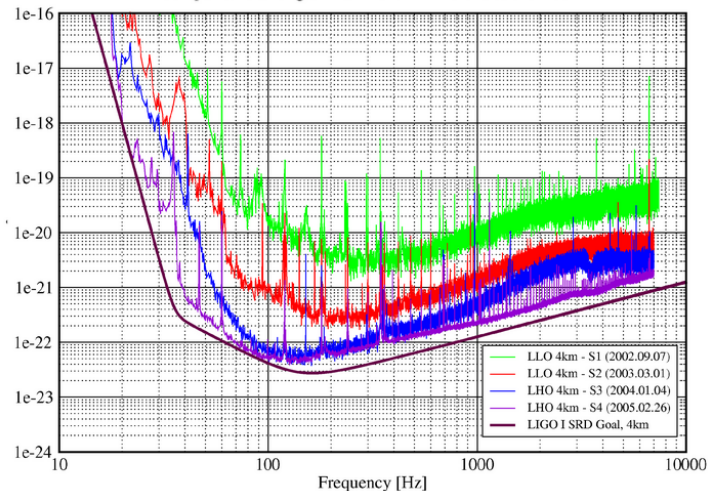
- electronics
- shot noise

# LIGO sensitivity, S1-S4 runs

## Best Strain Sensivities for the LIGO Interferometers

Comparisons among S1 - S4 Runs

LIGO-G050482-00-Z

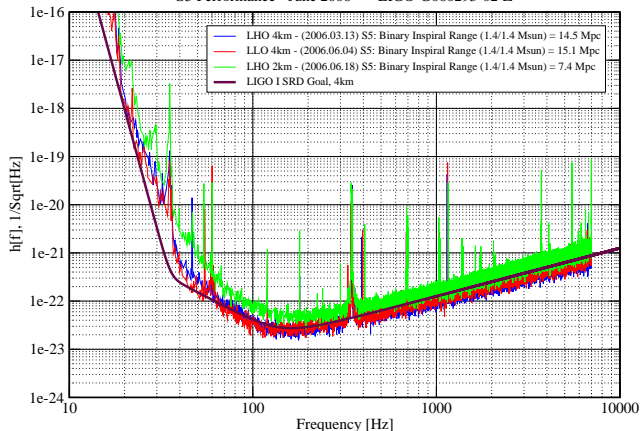


Inspiral search range during S4 was 8Mpc

# LIGO sensitivity, S5 run, June 2006

## Strain Sensitivity for the LIGO Interferometers

S5 Performance - June 2006 LIGO-G060293-02-Z



Inspiral search range during S5 is 14Mpc

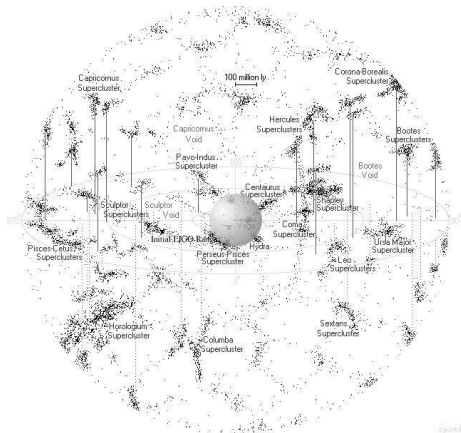
# Upgrade

## Goals

- Factor of 15 increase in sensitivity
- inspiral range from 20 Mpc to 350 Mpc
- Factor of 3000 in event rate  
**One day > entire 2-year initial data run**
- Quantum-noise-limited interferometer

## How

- better seismic isolation
- decreasing thermal noise
- higher laser power



# Seismic isolation





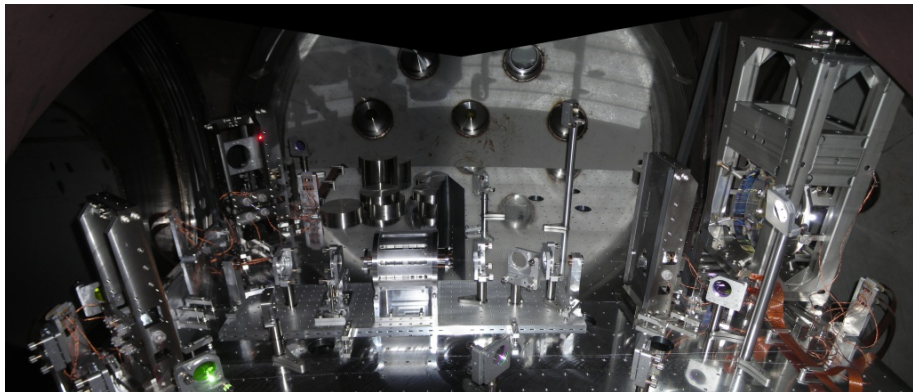
# Part of large system



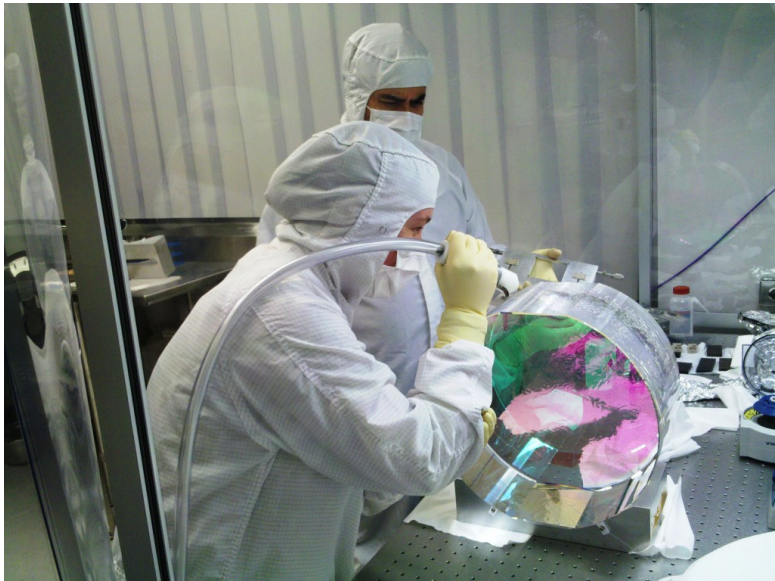
# Work in chamber



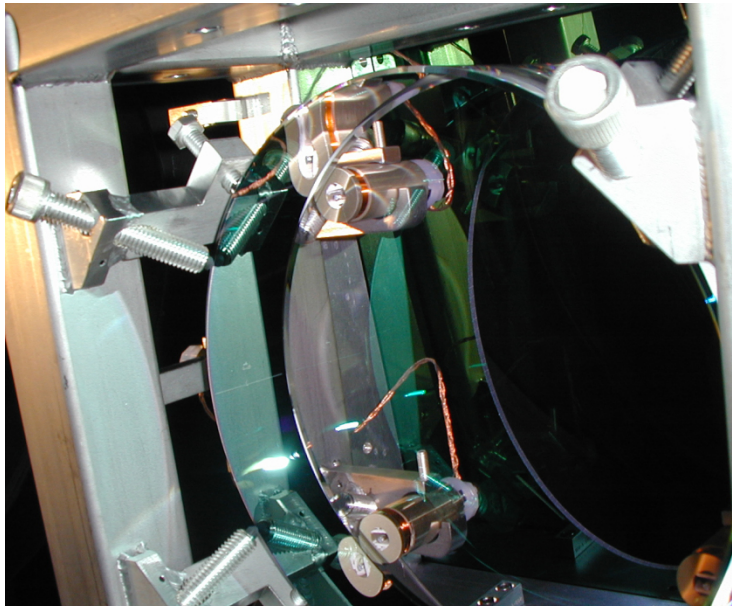
# Inside vacuum chamber



# Mirror

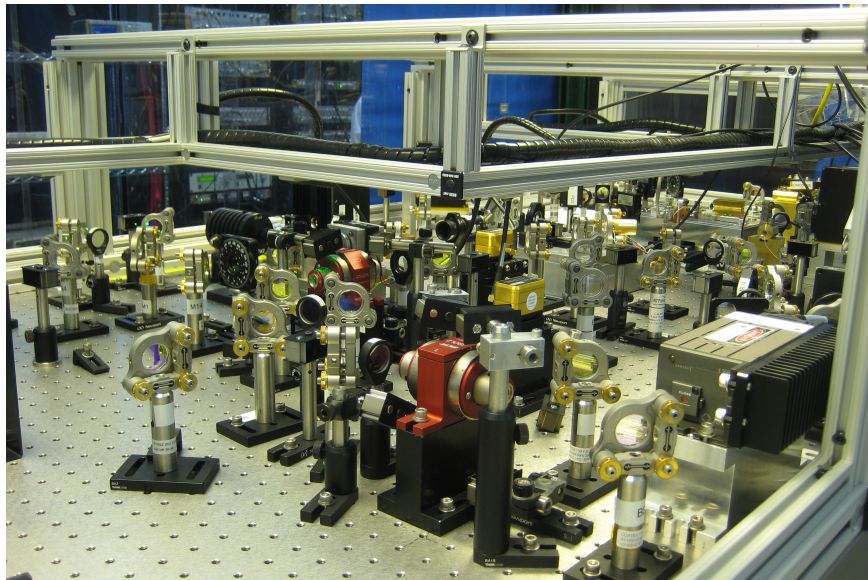


# Inner test mass





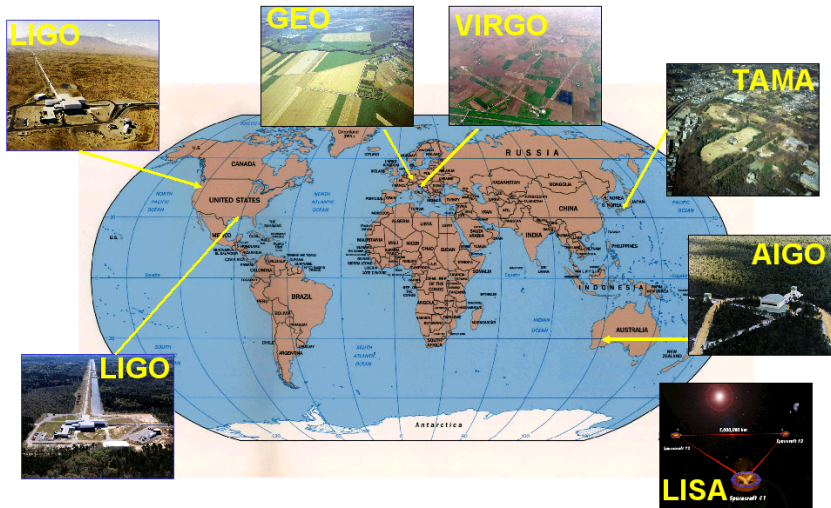
# Squeezer optical table



# Near by wild life



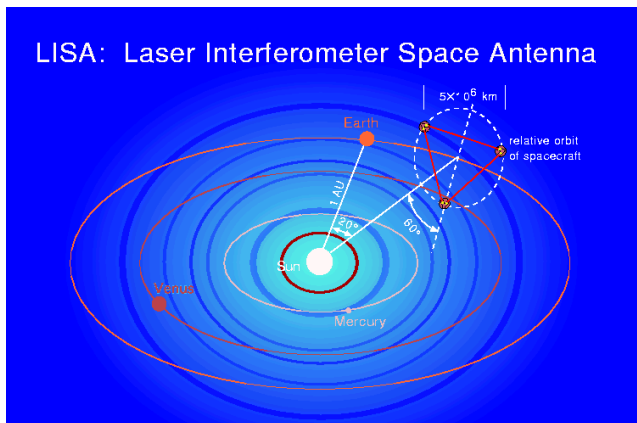
# World wide network of detectors





# Laser Interferometer Space Antenna (LISA)

- Three spacecraft in triangular formation
- separated by 5 million km
- Formation trails Earth by  $20^\circ$



# Additional links

[www.ligo.org](http://www.ligo.org)



You can help to detect a gravitational wave

[www.einsteinathome.org](http://www.einsteinathome.org)

